The SIM Astrometric Grid

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The SIM Astrometric Grid

Purpose

- Our Only Calibration of the Instrument
- Reference Mechanism for Science Measurements
 - Supplies reference for Science Object position
 - Used for determining the Baseline during data-taking

Content

- O(3000) Stable Stellar Objects distributed somewhat uniformly over the Celestial Sphere
- O(100) QSO's to establish Extragalactic Tie

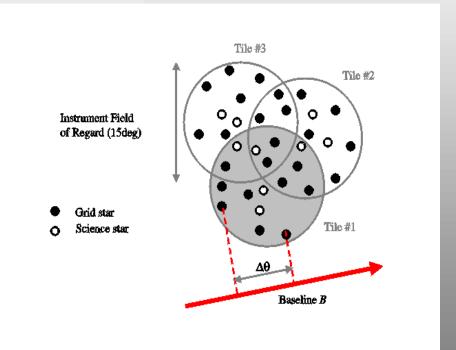
The Astrometric Grid

 SIM doesn't directly measure position or separation, all measurements are delays, one object at a time

$$d = s \bullet B + C + errors$$

- where:
 - d = delay measurement
 - s = object unit 3-vector
 - B = Baseline 3-vector
 - C = internal offset for inherent path length differences
- By bringing together many of these 1-d measurements, along with some a priori information on object positions, we can fit for the grid

The Grid Observations



- O(3000) Objects
- Individual Measurements are 1-d delays, not separations
- About 1/2 tile overlap
- ~1000 15° tiles per scan (solar exclusion of 45°)
- ~12-15 grid objects per tile

Developing the Astrometric Grid

- All Grid Objects repeatedly sampled during Mission to scan the whole sky ~4.5 times a Year.
- Object Measurements are done serially
 - Common Baseline Orientation during a single field-of-regard (which we call a "tile") ties delay measurements together for that tile.
 - Objects in tile overlap regions tie the tiles together for the 4π Grid
 - Celestial Sphere surveyed twice per Scan with Quasi-Orthogonal Baseline Projections to obtain Isotropic Position Errors.

The Astrometric Grid

 Thus, our observables are a series of delay measurements (and low-resolution baseline estimates) and a perturbed star catalog. From this, all we have to do is calculate a maximally consistent map of the universe.

Sequence of the Grid Simulation

- Generate O(3000) random truth objects
- Perturb truth objects to form knowledge objects
- Define grid scan campaign sequence
- Simulate delay measurement using truth catalog
- Use understanding of instrument to form Design Matrix
- Use knowledge catalog as our first a priori grid model
- Iterate a solution

The Design Matrix

- From the astrometric equation, we form a design matrix containing the details of the instrument performance as we understand it. This includes the way that the model parameters affect the results
 - Astrometric parameters α , δ , μ_{α} , μ_{β} , π
 - Baseline length and attitude
 - Effects we can model (focus of some current studies)
 - Stellar aberration
 - gravitational lensing from Sun, Jupiter

Solving the Grid

- Converge on a maximally consistent grid by iterating on
 - min|Ax b|

- where
 - · A is our design matrix
 - · x is a correction vector of the parameters
 - b is a vector of the differences between the observable delays:

$$b_i = d_i - (s_i \bullet B + C)_{model}$$

 Problem: The design matrix is large (100k x 100k) and sparse, so direct inverse is prohibitive for simulations and studies

Solving the Grid

 We exploit the sparseness of A by solving the associated set of normal equations:

$$(A^TA)x = A^Tb$$

- Find an approximate solution using the method of Conjugate Gradient on the Normal Equations. This approximation to the correction vector, **x** is then applied to the model and the procedure is reiterated
- Convergence usually occurs within 16-18 iterations
- This doesn't give a covariance matrix, but we can empirically approximate one when needed

Current Status of Grid Studies

- Simulation Fidelity
 - Ability to vary scan model (timing of tile observations)
 - Magnitude (fringe finding) uncertainties
 - Baseline metrology drifts enabled (including breaks)
 - Recently added multiple "pin objects"
- Completed Studies:
 - Metrology breaks between tiles (more robust than we thought)
- Current Studies:
 - Metrology Drifts within a tile:
 - Time dependent thermal
 - Angular dependent less than perfect optics
 - Non-Simple Grid Objects

Grid Object Selection

- Magnitude ~12
- Evenly distributed
- Simple Objects with linear Astrometric parameters
- Barycenter = Photocenter
- A range of distances
- Ability to study the object with ground-based measurements